



$$J = \frac{1}{2}$$

The following results are obtained using neutrinos associated with  $\tau^+$  or  $\tau^-$ . See Note on "Neutrino Mass," to be published in the PDG www update of November 2001.

Directly observed by the DONUT Collaboration (KODAMA 01). Existence indirectly established from  $\tau$  decay data combined with  $\nu$  reaction data. See for example FELDMAN 81. ALBRECHT 92Q rules out  $J = 3/2$  by establishing that the  $\rho^-$  is not in a pure  $H_{\rho}=-1$  helicity state in  $\tau^- \rightarrow \rho^- \nu_{\tau}$ .

### $\nu$ MASS

In the context of some models, it is possible that this weighted sum over mass eigenstates is the same as for the neutrinos produced in  $\mu$  decay.

In some of the ASTR and COSM papers listed below, the authors did not distinguish between weak and mass eigenstates.

VALUE (MeV)	CL%	EVTS	DOCUMENT ID	TECN	COMMENT
<b>&lt; 18.2</b>	95		<sup>1</sup> BARATE	98F ALEP	1991–1995 LEP runs
● ● ● We do not use the following data for averages, fits, limits, etc. ● ● ●					
< 28	95		<sup>2</sup> ATHANAS	00 CLEO	$E_{cm}^{ee} = 10.6$ GeV
< 27.6	95		<sup>3</sup> ACKERSTAFF	98T OPAL	1990–1995 LEP runs
< 30	95	473	<sup>4</sup> AMMAR	98 CLEO	$E_{cm}^{ee} = 10.6$ GeV
< 60	95		<sup>5</sup> ANASTASSOV	97 CLEO	$E_{cm}^{ee} = 10.6$ GeV
< 0.37 or >22			<sup>6</sup> FIELDS	97 COSM	Nucleosynthesis
< 68	95		<sup>7</sup> SWAIN	97 THEO	$m_{\tau}, \tau_{\tau}, \tau$ partial widths
< 29.9	95		<sup>8</sup> ALEXANDER	96M OPAL	1990–1994 LEP runs
<149			<sup>9</sup> BOTTINO	96 THEO	$\pi, \mu, \tau$ leptonic decays
<1 or >25			<sup>10</sup> HANNESTAD	96C COSM	Nucleosynthesis
< 71	95		<sup>11</sup> SOBIE	96 THEO	$m_{\tau}, \tau_{\tau}, B(\tau^- \rightarrow e^- \bar{\nu}_e \nu_{\tau})$
< 24	95	25	<sup>12</sup> BUSKULIC	95H ALEP	1991–1993 LEP runs
< 0.19			<sup>13</sup> DOLGOV	95 COSM	Nucleosynthesis
< 3			<sup>14</sup> SIGL	95 ASTR	SN 1987A
< 0.4 or > 30			<sup>15</sup> DODELSON	94 COSM	Nucleosynthesis
< 0.1 or > 50			<sup>16</sup> KAWASAKI	94 COSM	Nucleosynthesis
155–225			<sup>17</sup> PERES	94 THEO	$\pi, K, \mu, \tau$ weak decays
< 32.6	95	113	<sup>18</sup> CINABRO	93 CLEO	$E_{cm}^{ee} \approx 10.6$ GeV
< 0.3 or > 35			<sup>19</sup> DOLGOV	93 COSM	Nucleosynthesis
< 0.74			<sup>20</sup> ENQVIST	93 COSM	Nucleosynthesis
< 31	95	19	<sup>21</sup> ALBRECHT	92M ARG	$E_{cm}^{ee} = 9.4-10.6$ GeV
< 0.3			<sup>22</sup> FULLER	91 COSM	Nucleosynthesis
< 0.5 or > 25			<sup>23</sup> KOLB	91 COSM	Nucleosynthesis
< 0.42			<sup>22</sup> LAM	91 COSM	Nucleosynthesis

- <sup>1</sup> BARATE 98F result based on kinematics of 2939  $\tau^- \rightarrow 2\pi^- \pi^+ \nu_\tau$  and 52  $\tau^- \rightarrow 3\pi^- 2\pi^+ (\pi^0) \nu_\tau$  decays. If possible 2.5% excited  $a_1$  decay is included in 3-prong sample analysis, limit increases to 19.2 MeV.
- <sup>2</sup> ATHANAS 00 bound comes from analysis of  $\tau^- \rightarrow \pi^- \pi^+ \pi^- \pi^0 \nu_\tau$  decays.
- <sup>3</sup> ACKERSTAFF 98T use  $\tau \rightarrow 5\pi^\pm \nu_\tau$  decays to obtain a limit of 43.2 MeV (95%CL). They combine this with ALEXANDER 96M value using  $\tau \rightarrow 3h^\pm \nu_\tau$  decays to obtain quoted limit.
- <sup>4</sup> AMMAR 98 limit comes from analysis of  $\tau^- \rightarrow 3\pi^- 2\pi^+ \nu_\tau$  and  $\tau^- \rightarrow 2\pi^- \pi^+ 2\pi^0 \nu_\tau$  decay modes.
- <sup>5</sup> ANASTASSOV 97 derive limit by comparing their  $m_\tau$  measurement (which depends on  $m_{\nu_\tau}$ ) to BAI 96  $m_\tau$  threshold measurement.
- <sup>6</sup> FIELDS 97 limit for a Dirac neutrino. For a Majorana neutrino the mass region  $< 0.93$  or  $> 31$  MeV is excluded. These bounds assume  $N_\nu < 4$  from nucleosynthesis; a wider excluded region occurs with a smaller  $N_\nu$  upper limit.
- <sup>7</sup> SWAIN 97 derive their limit from the Standard Model relationships between the tau mass, lifetime, branching fractions for  $\tau^- \rightarrow e^- \bar{\nu}_e \nu_\tau$ ,  $\tau^- \rightarrow \mu^- \bar{\nu}_\mu \nu_\tau$ ,  $\tau^- \rightarrow \pi^- \nu_\tau$ , and  $\tau^- \rightarrow K^- \nu_\tau$ , and the muon mass and lifetime by assuming lepton universality and using world average values. Limit is reduced to 48 MeV when the CLEO  $\tau$  mass measurement (BALEST 93) is included; see CLEO's more recent  $m_{\nu_\tau}$  limit (ANASTASSOV 97).  
Consideration of mixing with a fourth generation heavy neutrino yields  $\sin^2 \theta_L < 0.016$  (95%CL).
- <sup>8</sup> ALEXANDER 96M bound comes from analyses of  $\tau^- \rightarrow 3\pi^- 2\pi^+ \nu_\tau$  and  $\tau^- \rightarrow h^- h^- h^+ \nu_\tau$  decays.
- <sup>9</sup> BOTTINO 96 assumes three generations of neutrinos with mixing, finds consistency with massless neutrinos with no mixing based on 1995 data for masses, lifetimes, and leptonic partial widths.
- <sup>10</sup> HANNESTAD 96C limit is on the mass of a Majorana neutrino. This bound assumes  $N_\nu < 4$  from nucleosynthesis. A wider excluded region occurs with a smaller  $N_\nu$  upper limit. This paper is the corrected version of HANNESTAD 96; see the erratum: HANNESTAD 96B.
- <sup>11</sup> SOBIE 96 derive their limit from the Standard Model relationship between the tau mass, lifetime, and leptonic branching fraction, and the muon mass and lifetime, by assuming lepton universality and using world average values.
- <sup>12</sup> BUSKULIC 95H bound comes from a two-dimensional fit of the visible energy and invariant mass distribution of  $\tau \rightarrow 5\pi(\pi^0) \nu_\tau$  decays. Replaced by BARATE 98F.
- <sup>13</sup> DOLGOV 95 removes earlier assumptions (DOLGOV 93) about thermal equilibrium below  $T_{\text{QCD}}$  for wrong-helicity Dirac neutrinos (ENQVIST 93, FULLER 91) to set more stringent limits. DOLGOV 96 argues that a possible window near 20 MeV is excluded.
- <sup>14</sup> SIGL 95 exclude massive Dirac or Majorana neutrinos with lifetimes between  $10^{-3}$  and  $10^8$  seconds if the decay products are predominantly  $\gamma$  or  $e^+ e^-$ .
- <sup>15</sup> DODELSON 94 calculate constraints on  $\nu_\tau$  mass and lifetime from nucleosynthesis for 4 generic decay modes. Limits depend strongly on decay mode. Quoted limit is valid for all decay modes of Majorana neutrinos with lifetime greater than about 300 s. For Dirac neutrinos limits change to  $< 0.3$  or  $> 33$ .
- <sup>16</sup> KAWASAKI 94 excluded region is for Majorana neutrino with lifetime  $> 1000$  s. Other limits are given as a function of  $\nu_\tau$  lifetime for decays of the type  $\nu_\tau \rightarrow \nu_\mu \phi$  where  $\phi$  is a Nambu-Goldstone boson.
- <sup>17</sup> PERES 94 used PDG 92 values for parameters to obtain a value consistent with mixing. Reexamination by BOTTINO 96 which included radiative corrections and 1995 PDG parameters resulted in two allowed regions,  $m_3 < 70$  MeV and  $140 \text{ MeV} < m_3 < 149$  MeV.
- <sup>18</sup> CINABRO 93 bound comes from analysis of  $\tau^- \rightarrow 3\pi^- 2\pi^+ \nu_\tau$  and  $\tau^- \rightarrow 2\pi^- \pi^+ 2\pi^0 \nu_\tau$  decay modes.

- <sup>19</sup> DOLGOV 93 assumes neutrino lifetime  $>100$  s. For Majorana neutrinos, the low mass limit is 0.5 MeV. KAWANO 92 points out that these bounds can be overcome for a Dirac neutrino if it possesses a magnetic moment. See also DOLGOV 96.
- <sup>20</sup> ENQVIST 93 bases limit on the fact that thermalized wrong-helicity Dirac neutrinos would speed up expansion of early universe, thus reducing the primordial abundance. FULLER 91 exploits the same mechanism but in the older calculation obtains a larger production rate for these states, and hence a lower limit. Neutrino lifetime assumed to exceed nucleosynthesis time,  $\sim 1$  s.
- <sup>21</sup> ALBRECHT 92M reports measurement of a slightly lower  $\tau$  mass, which has the effect of reducing the  $\nu_\tau$  mass reported in ALBRECHT 88B. Bound is from analysis of  $\tau^- \rightarrow 3\pi^- 2\pi^+ \nu_\tau$  mode.
- <sup>22</sup> Assumes neutrino lifetime  $>1$  s. For Dirac neutrinos. See also ENQVIST 93.
- <sup>23</sup> KOLB 91 exclusion region is for Dirac neutrino with lifetime  $>1$  s; other limits are given.

## $\nu$ (MEAN LIFE) / MASS

Measures  $[\sum |U_{\ell j}|^2 \Gamma_j m_j]^{-1}$ , where the sum is over mass eigenstates which cannot be resolved experimentally. Most of these limits apply to any  $\nu$  within the indicated mass range.

<u>VALUE (s/eV)</u>	<u>DOCUMENT ID</u>	<u>TECN</u>	<u>COMMENT</u>
● ● ● We do not use the following data for averages, fits, limits, etc. ● ● ●			
	24 DOLGOV	99 COSM	
	25 BILLER	98 ASTR	$m_\nu = 0.05\text{--}1$ eV
$>1 \times 10^{14}$	26 SIGL	95 ASTR	$m_\nu >$ few MeV
$>2.8 \times 10^{15}$	27,28 BLUDMAN	92 ASTR	$m_\nu <$ 50 eV
$< 10^{-12}$ or $> 5 \times 10^4$	29 DODELSON	92 ASTR	$m_\nu = 1\text{--}300$ keV
	30 GRANEK	91 COSM	Decaying $L^0$
	31 WALKER	90 ASTR	$m_\nu = 0.03 - \sim 2$ MeV
$>6.3 \times 10^{15}$	28,32 CHUPP	89 ASTR	$m_\nu <$ 20 eV
$>1.7 \times 10^{15}$	28 KOLB	89 ASTR	$m_\nu <$ 20 eV
	33 TERASAWA	88 COSM	$m_\mu = 30\text{--}70$ MeV
	34 KAWASAKI	86 COSM	$m_\nu > 10$ MeV
	35 LINDLEY	85 COSM	$m_\nu > 10$ MeV
	36 BINETRUY	84 COSM	$m_\nu \sim 1$ MeV
	37 SARKAR	84 COSM	$m_\nu = 10\text{--}100$ MeV
	38 HENRY	81 ASTR	$m_\nu = 16\text{--}20$ eV
	39 KIMBLE	81 ASTR	$m_\nu = 10\text{--}100$ eV
	40 REPHAELI	81 ASTR	$m_\nu = 30\text{--}150$ eV
	41 DERUJULA	80 ASTR	$m_\nu = 10\text{--}100$ eV
$>2 \times 10^{21}$	42 STECKER	80 ASTR	$m_\nu = 10\text{--}100$ eV
	43 DICUS	78 COSM	$m_\nu = 0.5\text{--}30$ MeV
$<3 \times 10^{-11}$	44 FALK	78 ASTR	$m_\nu < 10$ MeV
	45 COWSIK	77 ASTR	

- <sup>24</sup> DOLGOV 99 places limits in the (Majorana)  $\tau$ -associated  $\nu$  mass-lifetime plane based on nucleosynthesis. Results would be considerably modified if neutrino oscillations exist.
- <sup>25</sup> BILLER 98 use the observed TeV  $\gamma$ -ray spectra to set limits on the mean life of a radiatively decaying neutrino between 0.05 and 1 eV. Curve shows  $\tau_\nu/B_\gamma > 0.15 \times 10^{21}$  s at 0.05 eV,  $> 1.2 \times 10^{21}$  s at 0.17 eV,  $> 3 \times 10^{21}$  s at 1 eV, where  $B_\gamma$  is the branching ratio to photons.

- 26 SIGL 95 exclude  $1\text{ s} \lesssim \tau \lesssim 10^8\text{ s}$  for MeV-mass  $\tau$  neutrinos from SN 1987A decaying radiatively, and eliminates the lower limit using other published results.
- 27 BLUDMAN 92 sets additional limits by this method for higher mass ranges. Cosmological limits are also obtained.
- 28 Nonobservation of  $\gamma$ 's in coincidence with  $\nu$ 's from SN 1987A. Results should be divided by the  $\nu \rightarrow \gamma X$  branching ratio.
- 29 DODELSON 92 range is for wrong-helicity keV mass Dirac  $\nu$ 's from the core of neutron star in SN 1987A decaying to  $\nu$ 's that would have interacted in KAM2 or IMB detectors.
- 30 GRANEK 91 considers heavy neutrino decays to  $\gamma\nu_L$  and  $3\nu_L$ , where  $m_{\nu_L} < 100\text{ keV}$ . Lifetime is calculated as a function of heavy neutrino mass, branching ratio into  $\gamma\nu_L$ , and  $m_{\nu_L}$ .
- 31 WALKER 90 uses SN 1987A  $\gamma$  flux limits after 289 days to find  $\langle m/\tau \rangle > 1.1 \times 10^{15}\text{ eV s}$ .
- 32 CHUPP 89 should be multiplied by a branching ratio (about 1) and a detection efficiency (about 1/4), and pertains to radiative decay of any neutrino to a lighter or sterile neutrino.
- 33 TERASAWA 88 finds only  $10^2 < \tau < 10^4$  allowed for 30–70 MeV  $\nu$ 's from primordial nucleosynthesis.
- 34 KAWASAKI 86 concludes that light elements in primordial nucleosynthesis would be destroyed by radiative decay of neutrinos with  $10\text{ MeV} < m_\nu < 1\text{ GeV}$  unless  $\tau \lesssim 10^4\text{ s}$ .
- 35 LINDLEY 85 considers destruction of cosmologically-produced light elements, and finds  $\tau < 2 \times 10^3\text{ s}$  for  $10\text{ MeV} < m_\nu < 100\text{ MeV}$ . See also LINDLEY 79.
- 36 BINETRUY 84 finds  $\tau < 10^8\text{ s}$  for neutrinos in a radiation-dominated universe.
- 37 SARKAR 84 finds  $\tau < 20\text{ s}$  at  $m_\nu = 10\text{ MeV}$ , with higher limits for other  $m_\nu$ , and claims that all masses between 1 MeV and 50 MeV are ruled out.
- 38 HENRY 81 uses UV flux from clusters of galaxies to find  $\tau > 1.1 \times 10^{25}\text{ s}$  for radiative decay.
- 39 KIMBLE 81 uses extreme UV flux limits to find  $\tau > 10^{22} - 10^{23}\text{ s}$ .
- 40 REPHAELI 81 consider  $\nu$  decay  $\gamma$  effect on neutral  $H$  in early universe; based on M31 HI concludes  $\tau > 10^{24}\text{ s}$ .
- 41 DERUJULA 80 finds  $\tau > 3 \times 10^{23}\text{ s}$  based on CDM neutrino decay contribution to UV background.
- 42 STECKER 80 limit based on UV background; result given is  $\tau > 4 \times 10^{22}\text{ s}$  at  $m = 20\text{ eV}$ .
- 43 DICUS 78 considers effect of  $\nu$  decay photons on light-element production, and finds lifetime must be less than "hours." See also DICUS 77.
- 44 FALK 78 finds lifetime constraints based on supernova energetics.
- 45 COWSIK 77 considers variety of scenarios. For neutrinos produced in the big bang, present limits on optical photon flux require  $\tau > 10^{23}\text{ s}$  for  $m_\nu \sim 1\text{ eV}$ . See also COWSIK 79 and GOLDMAN 79.
-

## $\nu$ MAGNETIC MOMENT

Must vanish for a purely chiral massless Dirac neutrino. A massive Dirac or Majorana neutrino can have a transition magnetic moment connecting one mass eigenstate to another one. The experimental limits below usually cannot distinguish between the true (diagonal, in mass) magnetic moment and a transition magnetic moment.

The value of the magnetic moment for the standard SU(2) $\times$ U(1) electroweak theory extended to include massive neutrinos (see FUJIKAWA 80) is  $\mu_\nu = 3eG_F m_\nu / (8\pi^2 \sqrt{2}) = (3.20 \times 10^{-19}) m_\nu \mu_B$  where  $m_\nu$  is in eV and  $\mu_B = e\hbar/2m_e$  is the Bohr magneton. Given the upper bound  $m_\nu < 18$  MeV, it follows that for the extended standard electroweak theory,  $\mu_\nu < 6 \times 10^{-12} \mu_B$ .

Most of the astrophysical limits pertain to any neutrino.

VALUE ( $\mu_B$ )	CL%	DOCUMENT ID	TECN	COMMENT
<b><math>&lt;5.4 \times 10^{-7}</math></b>	90	46 COOPER-...	92 BEBC	$\nu_\tau e^- \rightarrow \nu_\tau e^-$
● ● ● We do not use the following data for averages, fits, limits, etc. ● ● ●				
$<8.0 \times 10^{-6}$	90	47 TANIMOTO	00 RVUE	$e^+ e^- \rightarrow \nu \bar{\nu} \gamma$
$<3 \times 10^{-12}$		48 RAFFELT	99 ASTR	Red giant luminosity
$<4 \times 10^{-10}$		49 RAFFELT	99 ASTR	Solar cooling
$<4.4 \times 10^{-6}$	90	ABREU	97J DLPH	$e^+ e^- \rightarrow \nu \bar{\nu} \gamma$ at LEP
$<3.3 \times 10^{-6}$	90	50 ACCIARRI	97Q L3	$e^+ e^- \rightarrow \nu \bar{\nu} \gamma$ at LEP
$<6.2 \times 10^{-11}$		51 ELMFORS	97 COSM	Depolarization in early universe plasma
$<2.7 \times 10^{-6}$	95	52 ESCRIBANO	97 RVUE	$\Gamma(Z \rightarrow \nu \nu)$ at LEP
$<5.5 \times 10^{-6}$	90	GOULD	94 RVUE	$e^+ e^- \rightarrow \nu \bar{\nu} \gamma$ at LEP
$\gtrsim 10^{-8}$		53 KAWANO	92 ASTR	Primordial $^4\text{He}$ abundance
$<5.6 \times 10^{-6}$	90	DESHPANDE	91 RVUE	$e^+ e^- \rightarrow \nu \bar{\nu} \gamma$
$<2 \times 10^{-12}$		54 RAFFELT	90 ASTR	Red giant luminosity
$<1 \times 10^{-11}$		55 RAFFELT	89B ASTR	Cooling helium stars
$<4. \times 10^{-6}$	90	56 GROTCHE	88 RVUE	$e^+ e^- \rightarrow \nu \bar{\nu} \gamma$
$<1.1 \times 10^{-11}$		55,57 FUKUGITA	87 ASTR	Cooling helium stars
$<6 \times 10^{-14}$		58 NUSSINOV	87 ASTR	Cosmic EM backgrounds
$<8.5 \times 10^{-11}$		57 BEG	78 ASTR	Stellar plasmons
46 COOPER-SARKAR 92 assume $f_{D_s}/f_\pi = 2$ and $D_s, \bar{D}_s$ production cross section = $2.6 \mu\text{b}$ to calculate $\nu$ flux.				
47 TANIMOTO 00 combined $e^+ e^- \rightarrow \nu \bar{\nu} \gamma$ data from VENUS, TOPAZ, and AMY.				
48 RAFFELT 99 is an update of RAFFELT 90. This limit applies to all neutrino flavors which are light enough ( $< 5$ keV) to be emitted from globular-cluster red giants. This limit pertains equally to electric dipole moments and magnetic transition moments, and it applies to both Dirac and Majorana neutrinos.				
49 RAFFELT 99 is derived from the helioseismological limit on a new energy-loss channel of the Sun. This limit applies to all neutrino flavors which are light enough ( $<1$ keV) to be emitted from the Sun. This limit pertains equally to electric dipole and magnetic transition moments, and it applies to both Dirac and Majorana neutrinos.				
50 ACCIARRI 97Q result applies to both direct and transition magnetic moments and for $q^2=0$ .				

- 51 ELMFORS 97 calculate the rate of depolarization in a plasma for neutrinos with a magnetic moment and use the constraints from a big-bang nucleosynthesis on additional degrees of freedom.
- 52 Applies to absolute value of magnetic moment.
- 53 KAWANO 92 lower limit is that needed to circumvent  ${}^4\text{He}$  production if  $m_\nu$  is between 5 and  $\sim 30 \text{ MeV}/c^2$ .
- 54 RAFFELT 90 limit valid if  $m_\nu < 5 \text{ keV}$ . It applies for a diagonal magnetic moment of a Dirac neutrino, or for a transition magnetic moment of a Majorana neutrino. In the latter case, the same analysis gives  $< 1.4 \times 10^{-12}$ . Limit at 95%CL obtained from  $\delta M_C$ .
- 55 Significant dependence on details of stellar properties.
- 56 GROUCH 88 combined data from MAC, ASP, CELLO, and Mark J.
- 57 If  $m_\nu < 10 \text{ keV}$ .
- 58 For  $m_\nu = 8\text{--}200 \text{ eV}$ . NUSSINOV 87 examines transition magnetic moments for  $\nu_\tau \rightarrow \nu_e$  and obtain  $< 3 \times 10^{-15}$  for  $m_\nu < 16 \text{ eV}$  and  $< 6 \times 10^{-14}$  for  $m_\nu > 4 \text{ eV}$ .

### $\nu$ ELECTRIC DIPOLE MOMENT

<u>VALUE (ecm)</u>	<u>CL%</u>	<u>DOCUMENT ID</u>	<u>TECN</u>	<u>COMMENT</u>
<b><math>&lt; 5.2 \times 10^{-17}</math></b>	95	59 ESCRIBANO	97 RVUE	$\Gamma(Z \rightarrow \nu\nu)$ at LEP

- 59 Applies to absolute value of electric dipole moment.

### $\nu$ CHARGE

<u>VALUE (units: electron charge)</u>	<u>DOCUMENT ID</u>	<u>TECN</u>	<u>COMMENT</u>
$< 2 \times 10^{-14}$	60 RAFFELT	99 ASTR	Red giant luminosity
$< 6 \times 10^{-14}$	61 RAFFELT	99 ASTR	Solar cooling
$< 4 \times 10^{-4}$	62 BABU	94 RVUE	BEBC beam dump
$< 3 \times 10^{-4}$	63 DAVIDSON	91 RVUE	SLAC electron beam dump

- • • We do not use the following data for averages, fits, limits, etc. • • •
- 60 This RAFFELT 99 limit applies to all neutrino flavors which are light enough ( $< 5 \text{ keV}$ ) to be emitted from globular-cluster red giants.
- 61 This RAFFELT 99 limit is derived from the helioseismological limit on a new energy-loss channel of the Sun, and applies to all neutrino flavors which are light enough ( $< 1 \text{ keV}$ ) to be emitted from the sun.
- 62 BABU 94 use COOPER-SARKAR 92 limit on  $\nu$  magnetic moment to derive quoted result.
- 63 DAVIDSON 91 use data from early SLAC electron beam dump experiment to derive charge limit as a function of neutrino mass.

$\nu_\tau$  REFERENCES

KODAMA	01	PL B504 218	K. Kodama <i>et al.</i>	(FNAL-653 Collab.)
ATHANAS	00	PR D61 052002	M. Athanas <i>et al.</i>	(CLEO Collab.)
TANIMOTO	00	PL B478 1	N. Tanimoto <i>et al.</i>	
DOLGOV	99	NP B548 385	A.D. Dolgov <i>et al.</i>	
RAFFELT	99	PRPL 320 319	G.G. Raffelt	
ACKERSTAFF	98T	EPJ C5 229	K. Ackerstaff <i>et al.</i>	(OPAL Collab.)
AMMAR	98	PL B431 209	R. Ammar <i>et al.</i>	(CLEO Collab.)
BARATE	98F	EPJ C2 395	R. Barate <i>et al.</i>	(ALEPH Collab.)
BILLER	98	PRL 80 2992	S.D. Biller <i>et al.</i>	(WHIPPLE Collab.)
ABREU	97J	ZPHY C74 577	P. Abreu <i>et al.</i>	(DELPHI Collab.)
ACCIARRI	97Q	PL B412 201	M. Acciarri <i>et al.</i>	(L3 Collab.)
ANASTASSOV	97	PR D55 2559	A. Anastassov <i>et al.</i>	(CLEO Collab.)
Also	98B	PR D58 119903 (erratum)	A. Anastassov <i>et al.</i>	(CLEO Collab.)
ELMFORS	97	NP B503 3	P. Elmfors <i>et al.</i>	
ESCRIBANO	97	PL B395 369	R. Escribano, E. Masso	(BARC, PARIT)
FIELDS	97	ASP 6 169	B.D. Fields, K. Kainulainen, K.A. Olive	(NDAM+)
SWAIN	97	PR D55 R1	J. Swain, L. Taylor	(NEAS)
ALEXANDER	96M	ZPHY C72 231	G. Alexander <i>et al.</i>	(OPAL Collab.)
BAI	96	PR D53 20	J.Z. Bai <i>et al.</i>	(BES Collab.)
BOTTINO	96	PR D53 6361	A. Bottino <i>et al.</i>	
DOLGOV	96	PL B383 193	A.D. Dolgov, S. Pastor, J.W.F. Valle	(IFIC, VALE)
HANNESTAD	96	PRL 76 2848	S. Hannestad, J. Madsen	(AARH)
HANNESTAD	96B	PRL 77 5148 (erratum)	S. Hannestad, J. Madsen	(AARH)
HANNESTAD	96C	PR D54 7894	S. Hannestad, J. Madsen	(AARH)
SOBIE	96	ZPHY C70 383	R.J. Sobie, R.K. Keeler, I. Lawson	(VICT)
BUSKULIC	95H	PL B349 585	D. Buskulic <i>et al.</i>	(ALEPH Collab.)
DOLGOV	95	PR D51 4129	A.D. Dolgov, K. Kainulainen, I.Z. Rothstein	(MICH+)
SIGL	95	PR D51 1499	G. Sigl, M.S. Turner	(FNAL, EFI)
BABU	94	PL B321 140	K.S. Babu, T.M. Gould, I.Z. Rothstein	(BART+)
DODELSON	94	PR D49 5068	S. Dodelson, G. Gyuk, M.S. Turner	(FNAL, CHIC+)
GOULD	94	PL B333 545	T.M. Gould, I.Z. Rothstein	(JHU, MICH)
KAWASAKI	94	NP B419 105	M. Kawasaki <i>et al.</i>	(OSU)
PERES	94	PR D50 513	O.L.G. Peres, V. Pleitez, R. Zukanovich Funchal	
BALEST	93	PR D47 R3671	R. Balest <i>et al.</i>	(CLEO Collab.)
CINABRO	93	PRL 70 3700	D. Cinabro <i>et al.</i>	(CLEO Collab.)
DOLGOV	93	PRL 71 476	A.D. Dolgov, I.Z. Rothstein	(MICH)
ENQVIST	93	PL B301 376	K. Enqvist, H. Uibo	(NORD)
ALBRECHT	92M	PL B292 221	H. Albrecht <i>et al.</i>	(ARGUS Collab.)
ALBRECHT	92Q	ZPHY C56 339	H. Albrecht <i>et al.</i>	(ARGUS Collab.)
BLUDMAN	92	PR D45 4720	S.A. Bludman	(CFPA)
COOPER-...	92	PL B280 153	A.M. Cooper-Sarkar <i>et al.</i>	(BEBC WA66 Collab.)
DODELSON	92	PRL 68 2572	S. Dodelson, J.A. Frieman, M.S. Turner	(FNAL+)
KAWANO	92	PL B275 487	L.H. Kawano <i>et al.</i>	(CIT, UCSD, LLL+)
PDG	92	PR D45, 1 June, Part II	K. Hikasa <i>et al.</i>	(KEK, LBL, BOST+)
DAVIDSON	91	PR D43 2314	S. Davidson, B.A. Campbell, D. Bailey	(ALBE+)
DESHPANDE	91	PR D43 943	N.G. Deshpande, K.V.L. Sarma	(OREG, TATA)
FULLER	91	PR D43 3136	G.M. Fuller, R.A. Malaney	(UCSD)
GRANEK	91	IJMP A6 2387	H. Granek, B.H.J. McKellar	(MELB)
KOLB	91	PRL 67 533	E.W. Kolb <i>et al.</i>	(FNAL, CHIC)
LAM	91	PR D44 3345	W.P. Lam, K.W. Ng	(AST)
RAFFELT	90	PRL 64 2856	G.G. Raffelt	(MPIM)
WALKER	90	PR D41 689	T.P. Walker	(HARV)
CHUPP	89	PRL 62 505	E.L. Chupp, W.T. Vestrand, C. Reppin	(UNH, MPIM)
KOLB	89	PRL 62 509	E.W. Kolb, M.S. Turner	(CHIC, FNAL)
RAFFELT	89B	APJ 336 61	G. Raffelt, D. Dearborn, J. Silk	(UCB, LLL)
ALBRECHT	88B	PL B202 149	H. Albrecht <i>et al.</i>	(ARGUS Collab.)
GROTCH	88	ZPHY C39 553	H. Grotch, R.W. Robinett	(PSU)
TERASAWA	88	NP B302 697	N. Terasawa, M. Kawasaki, K. Sato	(TOKY)
FUKUGITA	87	PR D36 3817	M. Fukugita, S. Yazaki	(KYOTU, TOKY)
NUSSINOV	87	PR D36 2278	S. Nussinov, Y. Rephaeli	(TELA)
KAWASAKI	86	PL B178 71	M. Kawasaki, N. Terasawa, K. Sato	(TOKY)
LINDLEY	85	APJ 294 1	D. Lindley	(FNAL)
BINETRUY	84	PL 134B 174	P. Binetruy, G. Girardi, P. Salati	(LAPP)
SARKAR	84	PL 148B 347	S. Sarkar, A.M. Cooper	(OXF, CERN)
FELDMAN	81	SLAC-PUB-2839	G.J. Feldman	(SLAC, STAN)

Santa Cruz APS.

HENRY	81	PRL 47 618	R.C. Henry, P.D. Feldman	(JHU)
KIMBLE	81	PRL 46 80	R. Kimble, S. Bowyer, P. Jakobsen	(UCB)
REPHAELI	81	PL 106B 73	Y. Rephaeli, A.S. Szalay	(UCSB, CHIC)
DERUJULA	80	PRL 45 942	A. De Rujula, S.L. Glashow	(MIT, HARV)
FUJIKAWA	80	PRL 45 963	K. Fujikawa, R. Shrock	(STON)
STECKER	80	PRL 45 1460	F.W. Stecker	(NASA)
COWSIK	79	PR D19 2219	R. Cowsik	(TATA)
GOLDMAN	79	PR D19 2215	T. Goldman, G.J. Stephenson	(LASL)
LINDLEY	79	MNRAS 188 15P	D. Lindley	(SUSS)
BEG	78	PR D17 1395	M.A.B. Beg, W.J. Marciano, M. Ruderman	(ROCK+)
DICUS	78	PR D17 1529	D.A. Dicus <i>et al.</i>	(TEXA, VPI, STAN)
FALK	78	PL 79B 511	S.W. Falk, D.N. Schramm	(CHIC)
COWSIK	77	PRL 39 784	R. Cowsik	(MPIM, TATA)
DICUS	77	PRL 39 168	D.A. Dicus, E.W. Kolb, V.L. Teplitz	(TEXA, VPI)

---

### OTHER RELATED PAPERS

---

WEINSTEIN	93	ARNPS 43 457	A.J. Weinstein, R. Stroynowski	(CIT, SMU)
-----------	----	--------------	--------------------------------	------------

---